

BAT AGN spectroscopic survey–II. X-ray emission and high-ionization optical emission lines

Simon Berney,¹ Michael Koss,^{1★†} Benny Trakhtenbrot,^{1‡} Claudio Ricci,² Isabella Lamperti,¹ Kevin Schawinski,¹ Mislav Baloković,³ D. Michael Crenshaw,⁴ Travis Fischer,⁴ Neil Gehrels,⁵ Fiona Harrison,³ Yasuhiro Hashimoto,⁶ Kohei Ichikawa,⁷ Richard Mushotzky,⁸ Kyuseok Oh,¹ Daniel Stern,⁹ Ezequiel Treister,¹⁰ Yoshihiro Ueda,⁷ Sylvain Veilleux⁸ and Lisa Winter¹¹

¹*Institute for Astronomy, Department of Physics, ETH Zurich, Wolfgang-Pauli-Strasse 27, CH-8093 Zurich, Switzerland*

²*Instituto de Astrofísica, Facultad de Física, Pontificia Universidad Católica de Chile, Casilla 306, Santiago 22, Chile*

³*Cahill Center for Astronomy and Astrophysics, California Institute of Technology, Pasadena, CA 91125, USA*

⁴*Department of Physics and Astronomy, Georgia State University, Astronomy Offices, One Park Place South SE, Suite 700, Atlanta, GA 30303, USA*

⁵*NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA*

⁶*Department of Earth Sciences, National Taiwan Normal University, No. 88, Section 4, Tingzhou Rd, Wenshan District, Taipei 11677, Taiwan Republic of China*

⁷*Department of Astronomy, Kyoto University, Kyoto 606-8502, Japan*

⁸*Astronomy Department, University of Maryland, College Park, MD 20742, USA*

⁹*Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, MS 169-21, Pasadena, CA 91109, USA*

¹⁰*Departamento de Astronomía, Universidad de Concepción, Concepción, Chile*

¹¹*Atmospheric and Environmental Research, 131 Hartwell Ave No. 4, Casilla 160-C, Lexington, MA 02421, USA*

Accepted 2015 September 17. Received 2015 September 17; in original form 2015 June 27

ABSTRACT

We investigate the relationship between X-ray and optical line emission in 340 nearby ($z \simeq 0.04$) AGN selected above 10 keV using *Swift* BAT. We find a weak correlation between the extinction corrected [O III] and hard X-ray luminosity ($L_{[\text{O III}]}^{\text{int}} \propto L_{14-195}$) with a large scatter ($R_{\text{Pear}} = 0.64$, $\sigma = 0.62$ dex) and a similarly large scatter with the intrinsic 2–10 keV to [O III] luminosities ($R_{\text{Pear}} = 0.63$, $\sigma = 0.63$ dex). Correlations of the hard X-ray fluxes with the fluxes of high-ionization narrow lines ([O III], He II, [Ne III] and [Ne V]) are not significantly better than with the low-ionization lines (H α , [S II]). Factors like obscuration or physical slit size are not found to be a significant part of the large scatter. In contrast, the optical emission lines show much better correlations with each other ($\sigma = 0.3$ dex) than with the X-ray flux. The inherent large scatter questions the common usage of narrow emission lines as AGN bolometric luminosity indicators and suggests that other issues such as geometrical differences in the scattering of the ionized gas or long-term AGN variability are important.

Key words: black hole physics – galaxies: active – galaxies: nuclei – quasars: general.

1 INTRODUCTION

Based on the unification model of active galactic nuclei (AGN; Urry, Maraschi & Phinney 1991; Antonucci 1993; Urry & Padovani 1995), narrow optical emission lines are emitted from the narrow-line region (NLR). The NLR is ionized by the extreme UV-soft X-ray emission from the AGN’s central engine. Afterwards, the il-

luminated interstellar gas in the NLR radiatively cools down by emitting the narrow optical lines. Unlike this optical emission, X-rays are emitted from much closer to the supermassive black hole. The UV photons radiated by the accretion disc are inverse Compton scattered inside the corona and can reach energies up to a few hundreds of keV (e.g. Netzer 2013). Thus, if AGN are accreting at a constant rate and constantly ionizing the NLR, we would naively expect a tight correlation between the X-ray emission from the corona and the optical emission lines from the NLR. This is why both high-ionization optical emission lines and X-rays are thought to be reliable tracers of the AGN bolometric luminosity (L_{bol}) and of the black hole growth (Heckman et al. 2005).

* E-mail: mike.koss@phys.ethz.ch

† Ambizione Fellow.

‡ Zwicky Fellow.

The bolometric luminosity (L_{bol}) of unobscured AGN estimated from the spectral energy distribution (SED) fitting technique using optical, UV, and X-ray data shows a tight correlation with hard X-ray data alone (e.g. 14–195 keV with a scatter of 0.3 dex; Vasudevan et al. 2009) though X-ray surveys may not detect heavily obscured Compton-thick AGN ($N_{\text{H}} > 10^{24} \text{ cm}^{-2}$). Strong observational correlation has also been found between UV emission and X-ray emission (Steffen et al. 2006; Young, Elvis & Risaliti 2010). The use of [O III] $\lambda 5007$ to infer AGN bolometric luminosities has a long history (e.g. Heckman et al. 2004; Kauffmann & Heckman 2009; LaMassa et al. 2009). It has been widely used to estimate the accretion rates of AGN using bolometric corrections based on the X-ray emission, the 5100 Å monochromatic continuum or the mid-IR emission (e.g. Lamastra et al. 2009). However, several studies have noted that the strength of [O III] $\lambda 5007$ is modulated by factors including covering factor of the NLR gas and its density (e.g. Netzer & Laor 1993), and ionization parameter (e.g. Baskin & Laor 2005).

Previous studies have compared X-ray emission with optical emission lines strengths in hard X-ray selected samples. Heckman et al. (2005) used a sample of 47 local AGN ($z < 0.2$) from the *RXTE* slew survey to analyse the correlation between 3–20 keV flux and the [O III] strength (Revnivtsev et al. 2004; Sazonov & Revnivtsev 2004). Other works have also made use of the early results from the *Swift* BAT survey (9-month catalogue; Tueller et al. 2008) to analyse optical and X-ray emission from AGN. Meléndez et al. (2008) studied the correlation between the [O III] $\lambda 5007$ line strength and two X-ray bands (2–10 keV and 14–195 keV) for 40 low-redshift ($z < 0.08$) BAT AGN from the 9-month catalogue. Winter et al. (2010) made a similar analysis with a slightly larger sample (64 BAT AGN). Finally, an analysis of the full 9-month catalogue has recently been completed (≈ 100 BAT AGN; Ueda et al. 2015, submitted). Netzer et al. (2006) studied a sample of 52 obscured AGN at high redshift ($z \approx 1$) suggesting that [O III] emission is weaker relative to 2–10 keV emission in high- L_X sources.

The *Swift* BAT 70-month all-sky survey enables us to analyse the optical and X-ray properties of a large sample of nearby AGN. It has detected 823 AGN to a flux limit of $\approx 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ over the whole sky (Baumgartner et al. 2013). Due to its detection method in the ultrahard X-rays (14–195 keV), the *Swift* BAT survey is not affected by obscuration (below Compton-thick levels) nor is it contaminated by stellar emission. Therefore, it also includes ‘buried’ obscured AGN that may not have been selected by optical surveys (e.g. Ueda et al. 2007).

In this work, we use a sample of BAT AGN based on the 70-month catalogue with an optical spectroscopic sample that is significantly larger than previous studies (340 objects). Our spectroscopic analysis also includes higher ionization lines not used previously (i.e. [Ne V] $\lambda 3426$) to study correlations between optical line emission and X-ray emission. The sample is presented in Section 2 and our analysis methods are described in Section 3. We study the correlation between [O III] and X-ray emission in Section 4.1, X-ray and other high-ionization lines (i.e. [Ne III] $\lambda 3869$, He II $\lambda 4686$, and [Ne V] $\lambda 3426$) in Section 4.2, and the correlation between only optical emission lines in Section 4.3. Finally, a review of our results and a discussion is presented in Section 5.

2 SAMPLE

The optical data are taken from the BAT AGN Spectroscopic Survey (BASS) Data Release 1 (Koss et al., in preparation), which gathered 559 optical spectra from targeted spectroscopic campaigns on BAT sources and public optical surveys (SDSS and 6dF; Jones et al.

2009; Alam et al. 2015). This sample is the largest and the most complete catalogue of AGN optical spectra selected from the *Swift* hard X-ray survey. It contains 67.6 per cent of the total AGN detections in the *Swift* BAT 70-month catalogue and has a mean redshift of $z = 0.10$.

From this parent sample, we first remove the spectra coming from the 6dF survey because they lack flux calibration. We remove all blazars listed in the Roma blazar catalog (BZCAT) v5.0 (Massaro et al. 2009) since the X-ray emission in such sources is affected by differential beaming. We also apply two redshift cuts. The first one, $z > 0.01$, avoids very nearby star-forming galaxies (such as M82) and ensures that a significantly large portion of the ionized region within the galaxy is covered (i.e. $> 0.2 \text{ kpc}$). We also limit to $z < 0.4$ to avoid redder emission lines such as H α from being shifted outside the observed wavelength interval. Additionally, we remove 10 spectra due to calibration problems around the [O III] line. Finally, we exclude SDSS J155334.73+261441.4, a gravitationally lensed quasar (Ofek et al. 2007). The final sample totals 340 optical spectra with $\langle z \rangle = 0.05$.

We define an emission line detection when we reach an $\text{SNR} > 3$ of the line with respect to the noise of the adjacent continuum. Out of these 340 objects, we have 327 detections of [O III], 301 of H α , 281 of [N II], 278 of [S II], 265 of H β , 206 of [O I], 144 of [Ne III], 106 of He II and 51 of [Ne V]. The intrinsic 2–10 keV flux values and the column density measurements provided for 335/340 AGN are based on a homogeneous model fitting using the best available X-ray data with simultaneous modelling of the 0.2–10 keV band from *XMM-Newton*, *Chandra*, or *Swift* XRT and the 14–195 keV band from *Swift* BAT (details in Ricci et al., in preparation).

We use the AGN emission-line diagnostics (e.g. Baldwin, Phillips & Terlevich 1981; Veilleux & Osterbrock 1987) as updated by Kewley et al. (2006) to further classify our sample. We classify as low ionization nuclear emission-line region (LINER) any object that shows a LINER classification in either the [O I] $\lambda 6300$ or [S II] $\lambda \lambda 6717, 6731$ diagnostic. For the remaining AGN, we use the [N II] $\lambda 6583$ diagnostic, which is most sensitive to AGN, to differentiate between Seyfert AGN and H II/composite/ambiguous (17/340). We use the presence of a broad line in H α to differentiate between type 1 (217/340) and type 2 objects (116/340). A more detailed description of the reduction procedures of the optical spectra, sample selection, and overall properties of the optical spectroscopic sample can be found in the BASS-I paper (Koss et al., in preparation).

3 ANALYSIS OF OPTICAL SPECTRA

We fit our sample of optical spectra using an extensive spectroscopic analysis toolkit for astronomy, PySpecKit, which uses a Levenberg–Marquardt algorithm for fitting (Ginsburg & Mirocha 2011). We implement separate methods for fitting sources with narrow lines only and sources with broad lines. We fit the H β (4650–5050 Å), H α (6250–6770 Å), and [O II] (3300–4000 Å) spectral regions separately. For narrow-line sources, we first fit the spectrum with stellar templates using the PPXF software (Penalized PiXel Fitting; Cappellari & Emsellem 2004) to remove the galaxy continuum and stellar absorption features that can effect line measurements. All emission line fits were inspected by eye to verify proper fitting.

We adopt a power-law fit (first order) to model the continuum and Gaussian components to model the emission lines for different spectral regions. For the H α spectral region, we fit the [O I] $\lambda 6300$, [N II] $\lambda 6548$, H α , [N II] $\lambda 6583$, and [S II] $\lambda 6716$, and [S II] $\lambda 6731$ lines. The narrow-line widths are tied together with an allowed

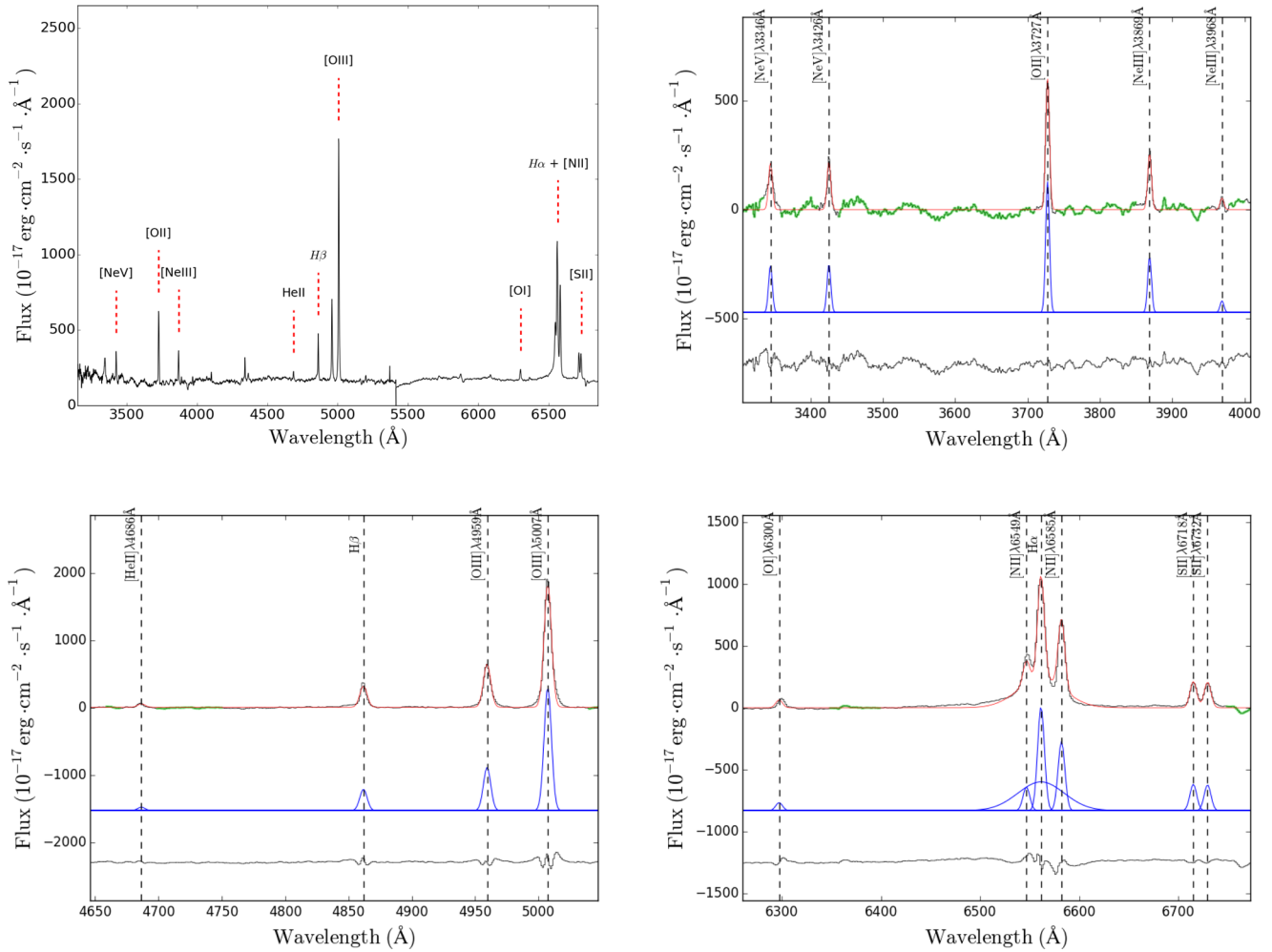


Figure 1. Example of the fitting procedure for MCG +06-16-28 observed with the Palomar Double Spectrograph. Top-left panel: spectrum of the object with all the fitted emission lines. Rest frame continuum subtracted fits are shown for the [O II] (top right), H β (bottom left), and H α (bottom right) spectral regions. The colours denote regions used for continuum (green), different Gaussian components of the fit (blue), and model fits (red) with residuals shown in the bottom panel.

variation of ± 8 Å and the systemic redshift is determined from the H α line. The relative strengths of [N II] $\lambda 6548$ and [N II] $\lambda 6583$ lines are fixed (at 1:2.94). For the H β spectral region, we fit the He II $\lambda 4686$, H β , [O III] $\lambda 4959$, and [O III] $\lambda 5007$ lines. The narrow-line widths are also tied together with an allowed variation of ± 8 Å and their central wavelengths are defined by the redshift of [O III] $\lambda 5007$. The intensity of [O III] $\lambda 4959$ relative to [O III] $\lambda 5007$ is fixed (at 1:2.86). For the [O II] spectral region, we fit the [Ne V] $\lambda 3346$, [Ne V] $\lambda 3426$, [O II] $\lambda 3727$, [Ne III] $\lambda 3869$, and [Ne III] $\lambda 3968$ lines.

Two Gaussian components are allowed for the H α and H β emission lines: a narrow component ($\text{FWHM} < 1000 \text{ km s}^{-1}$) and a broad component ($\text{FWHM} > 1000 \text{ km s}^{-1}$). The broad H β component is only allowed if a broad H α component is detected. In the case of a non-detection of the narrow H α or [S II] $\lambda \lambda 6717, 6731$ line (i.e. due to a very strong broad H α component, a noisy spectrum or bad sky subtraction), we use the FWHM of [O III] $\lambda 5007$ to constrain the widths of the narrow lines in the H α region.

To estimate the continuum for the H α complex, we use the wavelength regions 5800–6250 Å and 6750–7000 Å. For the H β complex, we use the 4660–4750 Å (except around He II) and 5040–5200 Å regions. The continuum around the [O II] spectral region is

usually more complicated to fit due to a non-linear shape or because it lies at the blue limit of the wavelength coverage. To fit this blue continuum, we use the region between 3300 and 4000 Å except where the emission lines are located (± 15 Å). To estimate errors in line fluxes and line widths, we use a Monte Carlo simulation that adds noise based on the error spectrum and reruns the fitting procedure 10 times. The flux uncertainty for the [O III] emission line is typically less than 1 per cent. To correct our line ratios for dust extinction, we use the narrow Balmer line ratio (H α /H β) assuming an intrinsic ratio of $R = 3.1$ (e.g. Ferland & Osterbrock 1986) and the Cardelli, Clayton & Mathis (1989) reddening curve when both the narrow H α and H β are detected. Fig. 1 shows an example spectrum and best model fit.

4 RESULTS

4.1 X-ray and [O III] $\lambda 5007$ flux

We first compare the X-ray flux (2–10 keV and 14–195 keV) to the [O III] $\lambda 5007$ emission line (Fig. 2). The relation between the [O III]-corrected flux and the BAT X-ray flux shows a scatter of $\sigma = 0.62$ dex and a Pearson correlation coefficient of $R_{\text{Pear}} = 0.35$. The

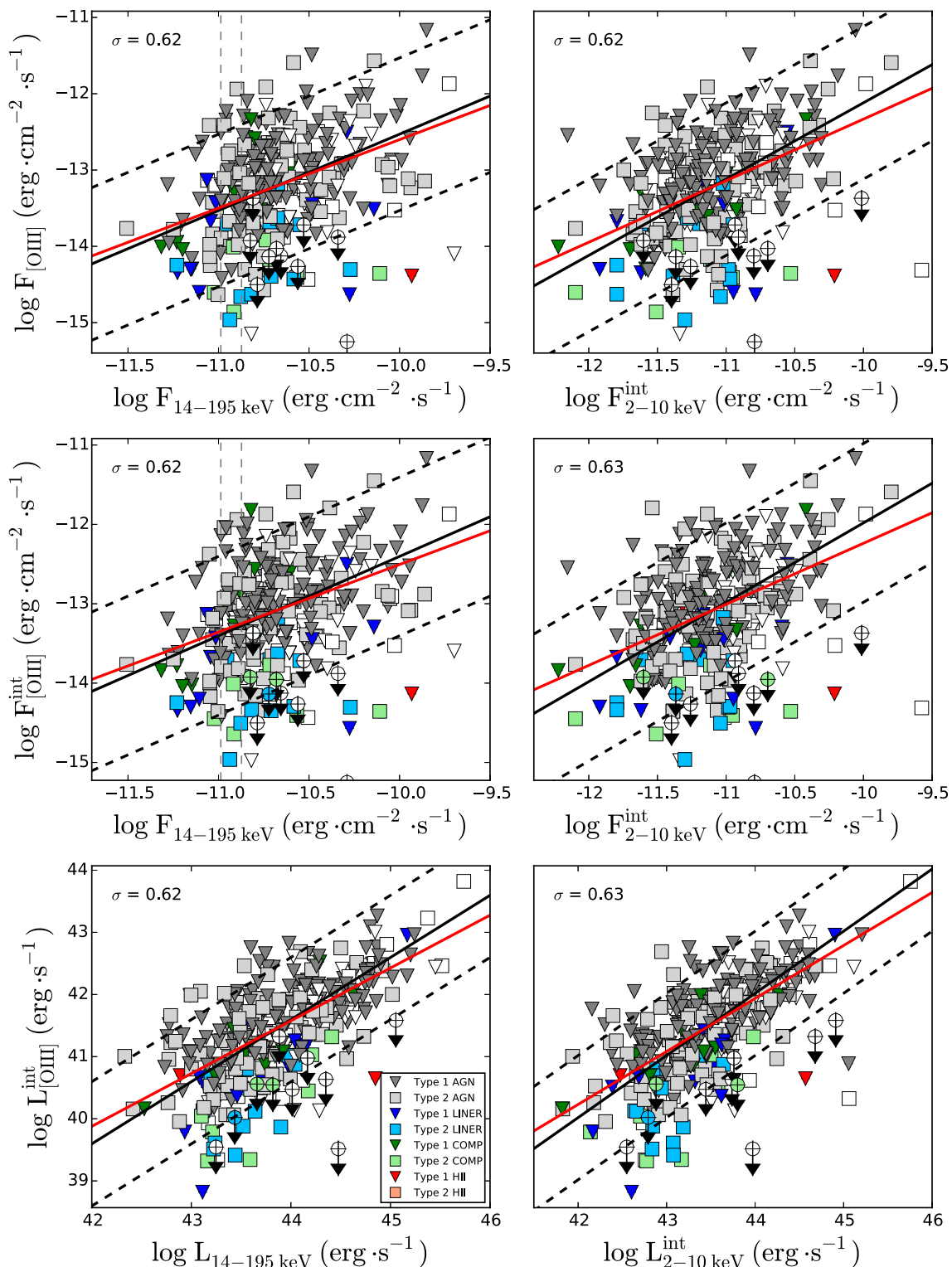


Figure 2. Plots of the relation of the [O III] $\lambda 5007$ emission line strength with the BAT X-ray emission (left) and with the 2–10 keV emission (right). The [O III]–X-ray relation exhibits a large scatter in all panels (≈ 0.6 dex). The black solid lines show a fit constrained by $Y \propto X$ with the dashed lines showing the ± 1 dex. The red lines illustrate the linear regression. For both line fits we removed the four largest outliers. The colour code comes from the [N II] BPT classification. White symbols do not have any BPT classification because of undetected emission lines. The AGN showing broad H α emission are classified as type 1. On the top and central left-hand panels, the grey dashed lines illustrate the 14–195 keV flux limits for 50 per cent (left line) and 90 per cent of the sky (right line). The upper limits were ignored in the analysis. Top panels: observed [O III] flux plotted against the X-ray fluxes. Central panels: extinction corrected [O III] flux plotted against the X-ray fluxes. Lower panels: extinction corrected [O III] luminosity versus the two X-ray luminosities.

Table 1. [O III]–X-ray luminosity relations for different sub-samples.

Y	X	Sample	N	a	b	σ	$\langle X \rangle$	$\langle Y \rangle$
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
$\log L_{[\text{O III}]}^{\text{int}}$	$\log L_{14-195}$	All	327	1.29 ± 0.05	-15 ± 2	0.61	43.9	41.4
		Type 1–1.9	215	1.17 ± 0.06	-9.6 ± 2.6	0.53	44.0	41.6
		Type 2.0	112	1.40 ± 0.10	-20 ± 4	0.69	43.7	41.1
		Seyferts	241	1.16 ± 0.05	-9.3 ± 2.4	0.53	43.9	41.5
		LINERs	17	1.8 ± 0.3	-39 ± 13	0.70	43.6	40.5
		H II/comp/ambiguous	17	1.2 ± 0.2	-11 ± 8	0.74	43.7	40.9
		All	321	1.23 ± 0.05	-12 ± 2	0.61	43.5	41.4
		Type 1–1.9	209	1.10 ± 0.06	-6.5 ± 2.5	0.56	43.6	41.6
		Type 2.0	112	1.30 ± 0.10	-15 ± 4	0.69	43.3	41.1
		Seyferts	238	1.13 ± 0.05	-7.5 ± 2.3	0.54	43.4	41.5
$\log L_{[\text{O III}]}^{\text{int}}$	$\log L_{2-10}^{\text{int}}$	LINERs	17	1.8 ± 0.3	-37 ± 12	0.62	43.1	40.5
		H II/comp/ambiguous	16	1.0 ± 0.2	-4 ± 8	0.62	43.2	40.8

Note. (1) variable Y ; (2) variable X ; (3) AGN type : type 1–1.9 are all the objects with broad H α detection, type 2 are all the objects without broad detection, LINERs are selected by the [S II] and [O I] diagrams, H II/comp/ambiguous and Seyferts are selected by the [N II] diagram (except for previously selected LINERs); (4) size of sample; (5) slope with its 1σ uncertainty; (6) regression intercept with its 1σ uncertainty. The equation of the linear regression is $Y = aX + b$; (7) standard deviation; (8) mean value of X ; (9) mean value of Y .

result is very similar when comparing [O III] to the 2–10 keV flux ($\sigma = 0.63$, $R_{\text{Pear}} = 0.38$). The Pearson p -value representing the probability of the hypothesis to have no correlation at all between the two X-ray fluxes and the [O III] flux is $<10^{-10}$ because of the large sample size. The correlation in luminosity space is better ($R_{\text{Pear}}^{14-195} = 0.64$, $R_{\text{Pear}}^{2-10} = 0.63$) than in flux space, as expected because the same multiplicative factor ($4\pi d_L^2$) is applied to both the X-ray flux and the [O III] flux to obtain the luminosities.¹ We finally compare the [O III] $\lambda 5007$ flux both with ($F_{[\text{O III}]}^{\text{int}}$) and without ($F_{[\text{O III}]}$) extinction correction to the two X-ray band fluxes. We find no significant improvement in the scatter of the different relations (~ 0.01 dex) using the Balmer decrement corrected [O III] fluxes.

Next, we compare a one-to-one relation and the linear regression in logarithmic space to fit our data. For both slope fits, we remove the four largest outliers. For the line fitting, we use the PYTHON routine `scipy.optimize.curve_fit` which provides a least squares approach (OLS($Y|X$)) with the Levenberg–Marquardt gradient method for convergence. The black solid lines in the panels of Fig. 2 show the best one-to-one relation (which would imply that $L_{[\text{O III}]}^{\text{int}} \propto L_{14-195}$). The best fit for the linear regression, shown with the red lines, is $L_{[\text{O III}]}^{\text{int}} \propto (L_{14-195})^{0.85 \pm 0.05}$ and $L_{[\text{O III}]}^{\text{int}} \propto (L_{2-10}^{\text{int}})^{0.85 \pm 0.05}$. To test if the linear regression is a significantly better fit, we run a F -test. The F -test compares the sum of the residuals squared of two different models. We obtain a F -value of 7.66 for the 14–195 keV flux and of 7.81 for the 2–10 keV flux. The resulting p -values are lower than 1 per cent suggesting that the linear regression between the [O III] luminosity and the two X-ray bands is a significantly better model than a one-to-one relation. However, the improvement in the scatter is very small (≈ 0.01 dex). Therefore, we decided to use the one-to-one relation (the simplest model) for the rest of the analysis.

We also use the *Astronomy Survival Analysis* (ASURV) package from the STATCODES suite of utilities developed by Eric Feigelson (Feigelson 1985). We determine the correlation between [O III] and X-ray luminosity in the presence of the 13 upper limits in [O III].

Specifically, we use the EM algorithm, though the Buckley James yields similar results. We find the best-fitting correlation is $L_{[\text{O III}]}^{\text{int}} \propto (L_{14-195})^{0.83 \pm 0.06}$ which is consistent with the least squares approach suggesting the small number of non-detections do not significantly change the slope of the correlation.

Finally, we use the ordinary least squared (OLS) bisector method of line fitting which is more reliable at recovering the intrinsic slope in the presence of uncertainties in both the X and Y variables (e.g. Isobe 1990). Table 1 lists the best-fitting linear regressions for the luminosity–luminosity plots using the OLS bisector method as well as the standard deviation for the complete sample and various sub-samples (Type 1–1.9, Type 2, Seyferts, LINERs, and H II/composite/ambiguous). The LINERs and H II/composite objects (from the BPT classification; Baldwin et al. 1981; Veilleux & Osterbrock 1987) have on average lower X-ray and [O III] luminosities than Seyfert galaxies: $\langle \log L_{[\text{O III}]}^{\text{LIN}}, \text{int} \rangle = 40.5$; $\langle \log L_{[\text{O III}]}^{\text{comp}}, \text{int} \rangle = 40.9$; $\langle \log L_{[\text{O III}]}^{\text{Sey}}, \text{int} \rangle = 41.5$.

4.1.1 Correlations with obscuration and luminosity

The two upper panels of Fig. 3 show the $L_{[\text{O III}]}^{\text{int}}/L_X$ ratio as a function of X-ray luminosity. The red dots show the average trend and the error bars the average 1σ scatter. First, we notice that there is no significant difference between the 2–10 keV (right-hand panels) and the BAT (left-hand panels) X-ray luminosity relations. There is a slight decrease of the [O III]/X-ray flux ratio when we go to higher X-ray luminosities, which is consistent though with being constant at a 3σ level. Fig. 4 illustrates the luminosity dependence of the flux ratio using histograms. The scatter slightly decreases towards high X-ray luminosity. This trend is not statistically significant, and is probably driven by the increase of the unobscured AGN fraction with high luminosities (receding torus model; e.g. Lawrence 1991; Simpson 1998, 2005; Brightman et al. 2015) since unobscured AGN exhibit lower scatter than the obscured AGN (see central panels of Fig. 3, see below). The two lower panels of Fig. 4 show a strong increase of the flux ratio as a function of the [O III] luminosity.

Dust and obscuration might still play a significant role in the relation between optical and X-ray emission even if the optical

¹ d_L varies here from 43 Mpc to 2.2 Gpc

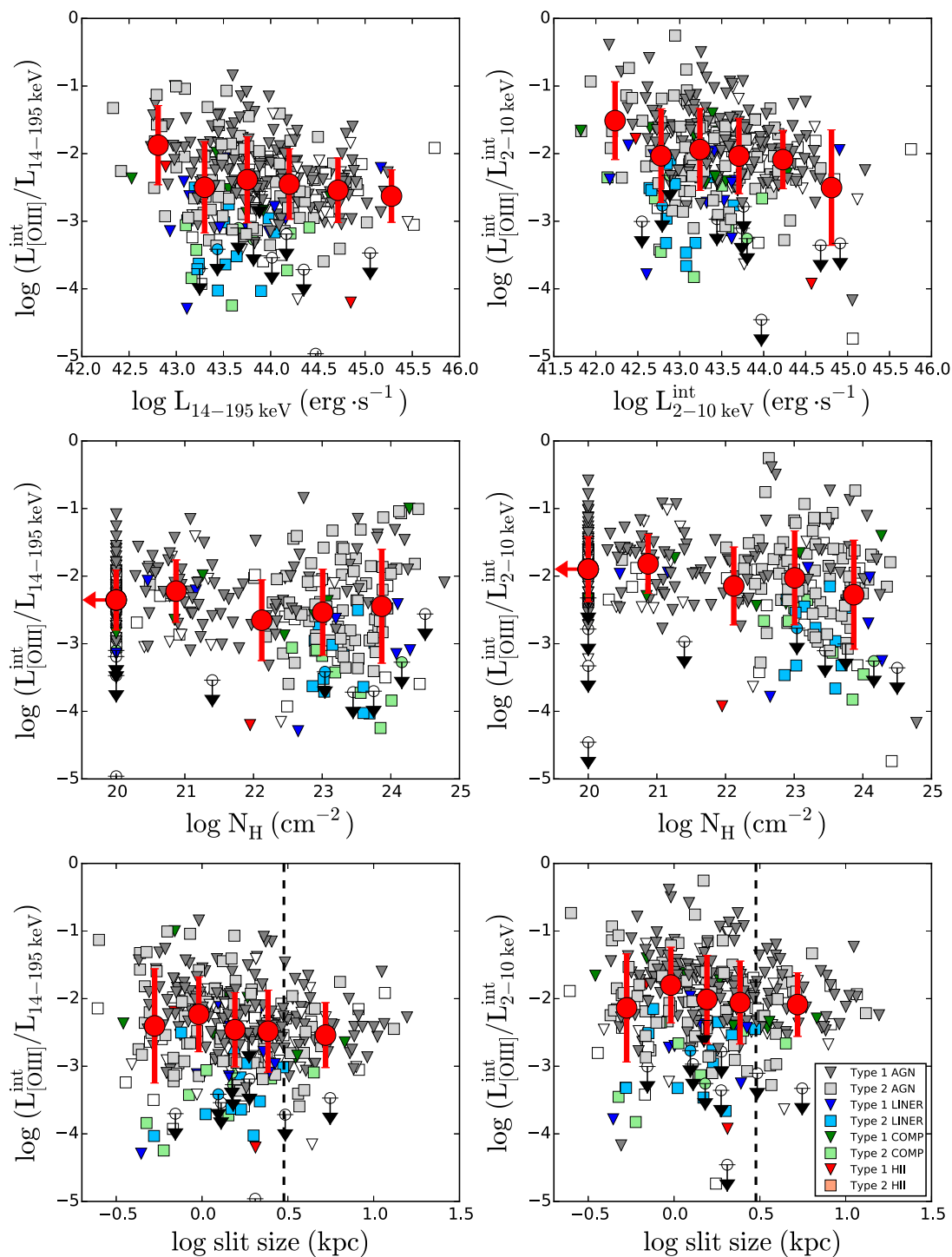


Figure 3. Plots of the $L_{\text{[O III]}}^{\text{int}}/L_X$ ratio as a function of X-ray luminosity (top panels), column density (central panels), and physical slit size (lower panels). The colour codes represent the [N II] BPT classification. White symbols do not have any BPT classification because of undetected emission lines. The red dots represent average binned values and the red bars show 1σ deviations in the bins. The first bin of the central panels represents an upper limit where the column density is consistent with the Galactic neutral hydrogen. The horizontal black lines on the lower panels illustrate the 3 kpc value. The scatter increases for $N_H > 10^{22} \text{ cm}^{-2}$ with no change in average $L_{\text{[O III]}}^{\text{int}}/L_X$. For slit size, the scatter decreases with increasing slit size because of a higher fraction of unobscured AGN at higher redshifts, but the average $L_{\text{[O III]}}^{\text{int}}/L_X$ is constant. The upper limits were ignored in the analysis.

emission lines are corrected for host galaxy extinction using the Balmer decrement. The obscuring torus scatters a portion of the ionizing light from the nucleus and therefore also has an indirect effect on the optical emission line intensities. The two central panels of Fig. 3 show the $L_{\text{[O III]}}^{\text{int}}/L_X$ ratio as a function of the integrated

column density N_H . Our results show that the trend of the luminosity ratio is consistent with being constant from very low to very high levels of obscuration. This is true for the intrinsic 2–10 keV emission as well as for the 14–195 keV band. We obtain average ratios of $\log L_{\text{[O III]}}^{\text{int}}/L_{14-195} = -2.42$ and $\log L_{\text{[O III]}}^{\text{int}}/L_{2-10} = -2.01$.

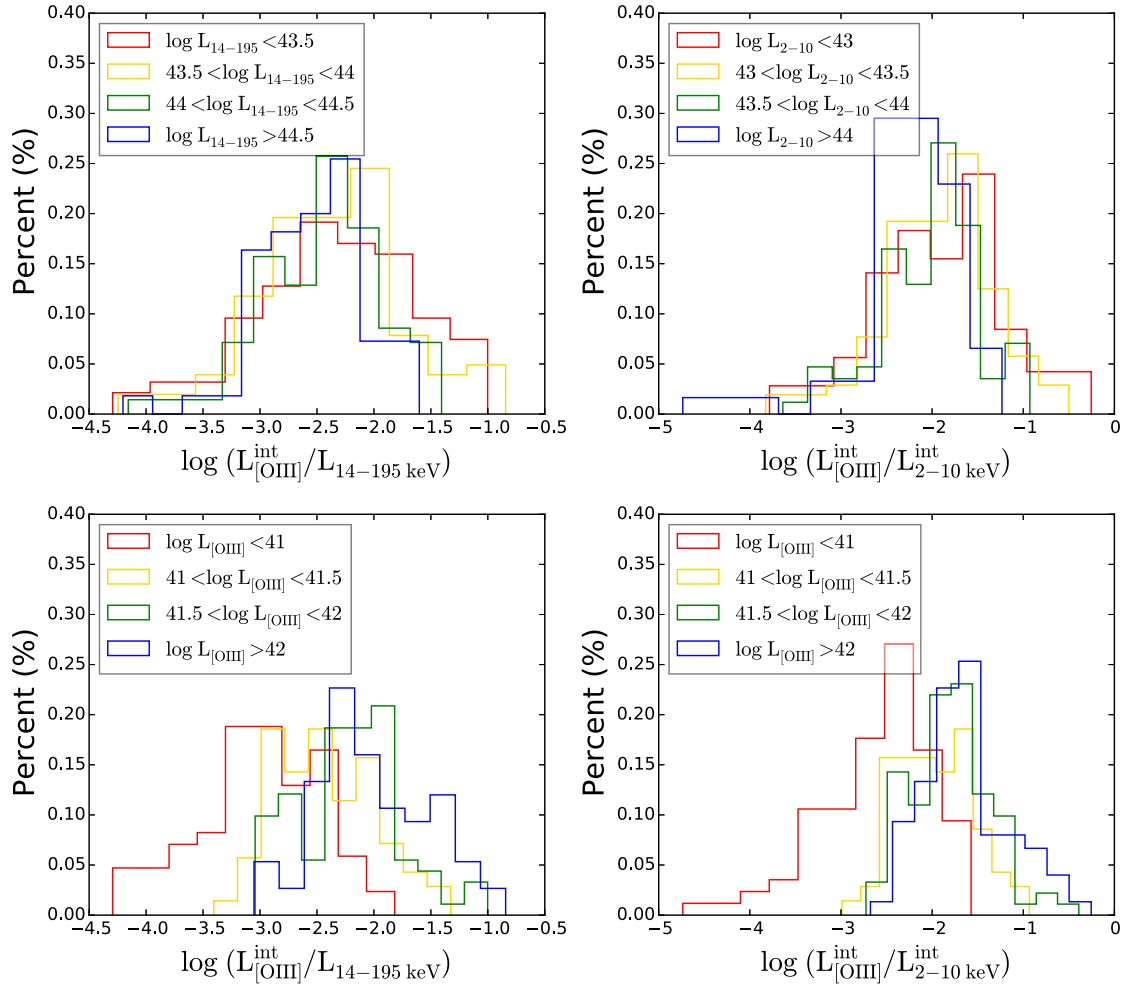


Figure 4. Histograms of the $L_{[\text{O III}]}^{\text{int}}/L_X$ ratio binned by X-ray luminosity (top panels) and [O III] luminosity (lower panels). The two lower panels show a strong positive correlation of the ratio with the [O III] luminosity.

However, the scatter strongly increases for $N_{\text{H}} > 10^{22} \text{ cm}^{-2}$. We observe a scatter of $\sigma = 0.47$ for low-obscuration AGN ($N_{\text{H}} < 10^{22} \text{ cm}^{-2}$) and of $\sigma = 0.72$ for high-obscuration AGN ($N_{\text{H}} > 10^{22} \text{ cm}^{-2}$) in the case of the BAT X-ray band. The result is similar for the 2–10 keV emission. In order to quantify this change in scatter, we run a Levene’s test. This test checks whether the two samples are drawn from distributions having a similar variance. We obtain a p -value = 6.52×10^{-6} for the BAT band and of 2.35×10^{-4} for the 2–10 keV band, implying a highly significant increase in the scatter from low to high obscuration in both X-ray bands.

4.1.2 Slit size and dilution

Due to different slit widths (0.75–3.0 arcsec) used to observe the BAT AGN and the various redshifts of these AGN ($0.01 < z < 0.4$), our spectroscopic data cover regions between 0.2 kpc to the size of the entire host galaxy. In addition to that, the NLR size may vary from one galaxy to another and the variation is strongly dependent on the power of the AGN (e.g. Mor, Netzer & Elitzur 2009; Hainline et al. 2013). This change in the size of the observed region and in the size of the NLR has various implications for our measurements. First, the region we probe with the spectroscopic observations might not cover the entire ionized [O III] region for luminous sources at low redshift. We might therefore underestimate the [O III] flux for

these objects. Alternatively in low-luminosity AGN, the region we probe with the slit could be much larger than the size of the NLR with a significant contaminating component from the host galaxy.

We study $L_{[\text{O III}]}^{\text{int}}/L_X$ as a function of the physical slit size in kpc (lower panels of Fig. 3) to study these issues. Our results show that the scatter slightly decreases towards larger slit sizes while the ratio $L_{[\text{O III}]}^{\text{int}}/L_X$ is constant. To further quantify this trend, we split our sample in two based on slit sizes > 3 kpc and slit sizes < 3 kpc. We observe that the scatter decreases by respectively 0.16 and 0.20 dex for the 14–195 keV and 2–10 keV band. However, this decrease in scatter is probably driven by the increase of the unobscured AGN fraction with high redshift rather than simply larger physical slit sizes.

4.1.3 Sample inhomogeneity

The BASS survey gathered optical spectra from various instruments with different data qualities (i.e. resolution, sensitivity, wavelength range). This lack of homogeneity in the data could artificially increase the scatter in the [O III]–X-ray relation. For this reason, we also analyse the SDSS sub-sample (110 AGN) because of their high quality in flux calibration (Fig. 5). We find that the scatter decreases by 0.09 dex for the 14–195 keV and 0.08 dex for the 2–10 keV band. In order to have a quantitative result for this change in scatter,

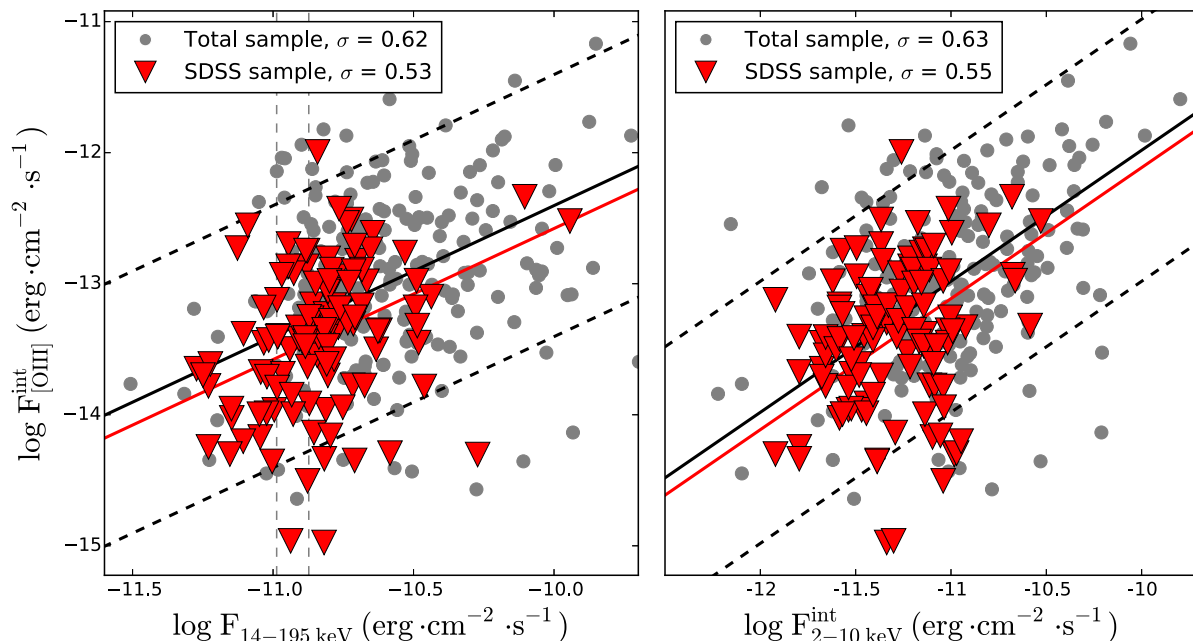


Figure 5. Plots of the [O III] $\lambda 5007$ extinction corrected flux versus the two X-ray fluxes. The red triangles represent the sources observed by SDSS (110 objects). The black line is a one-to-one fit of the total sample and the red line a one-to-one fit of the SDSS sub-sample. The scatter and the fit of the sub-sample are comparable to the ones of the main sample.

we run again a Levene's test. The resulting p -values for the SDSS sub-sample are, respectively, 0.062 and 0.35 for the BAT and 2–10 keV X-ray bands which is not statistically significant. The fit of the SDSS sub-sample (red line) is similar to the fit of the total sample (black line) for both X-ray bands. This implies that we do not have any statistical offset due to systematics (like flux calibration issue). We conclude that there is no significant difference between the total sample and the SDSS sub-sample.

4.2 X-ray correlation with high- and low-ionization lines

Lines produced from high ionization (>40 eV) have the advantage of being less contaminated by host galaxy emission than lines produced from low ionization (<20 eV). Moreover, these lines are emitted in less extended regions, which implies that we cover a larger portion of their emitting regions in our observations than for low-ionization lines. Finally, we include the H α and H β emission lines as these lines are either not or less affected with respect to other lines by variations in metal abundances, ionization levels, and collisional effects.

Fig. 6 illustrates the various emission line fluxes as a function of the two X-ray fluxes. Our data show that the correlation is slightly stronger with [Ne III], He II or [Ne V] than with [O III], H α or [S II], especially for the BAT X-ray band. Their scatter is smaller by about ~ 0.09 dex and the correlation coefficient increases from $R_{\text{Pear}} \sim 0.36$ to $R_{\text{Pear}} \sim 0.44$. However, the [Ne V] emission line is only detected in optically selected AGNs with very bright emission lines whereas the [O III] line is detected in many faint AGN such as the LINERs and composite objects which has the effect to increase the scatter. To overcome these possible biases, we need to compare the same samples together. This is what we show in Table 2. In this table, the reported standard deviations and Pearson coefficients are calculated based on a sub-sample of objects having both the detection of the considered line and an [O III] detection. In order to quantitatively investigate if the flux correlation of X-rays with [S II],

H α , [Ne III], He II or [Ne V] differ significantly from the one with [O III], we run a z -test based on the two Pearson correlation coefficients (Fisher r -to- z transformation). The p -value is also reported in Table 2. The result of this statistical analysis shows that we have p -value > 0.5 for all emission lines. The small differences between the linear relations are therefore not statistically significant.

4.3 Relationships between optical emission lines

Fig. 7 shows the relation between [O III] and other optical emission lines. We see from the scatter (~ 0.3 dex) that these emission lines are tightly correlated together. The mean flux ratios, standard deviations and Pearson coefficients are summarized in Table 3. One can especially notice the [O III]–[Ne III] relation that exhibits a scatter of $\sigma = 0.24$ dex. The two transitions originate from ions with comparable ionization energies (35.1 and 41 eV, respectively), and from gas with comparable critical densities (5×10^5 and $7 \times 10^6 \text{ cm}^{-3}$, respectively; Kogure & Leung 2007) which may explain the small scatter.

5 SUMMARY AND DISCUSSION

We have presented a study of emission line properties of the largest hard X-ray selected sample of AGN in the local Universe. Our conclusions are as follows.

- (i) We find very weak correlations with a large scatter (0.62–0.63 dex) between the two X-ray fluxes (2–10 keV and 14–195 keV band) and the observed and dust-corrected [O III]. The ratios between the X-ray fluxes and the [O III]-corrected flux are $F_{14-195} = (267 \pm 22) \times F_{[\text{O III}]}^{\text{int}}$ and $F_{2-10} = (102 \pm 8) \times F_{[\text{O III}]}^{\text{int}}$.
- (ii) The correlation of emission lines with high-ionization levels ([Ne III], He II, [Ne V]) with the X-ray flux is not significantly better than for [O III]. This is also true for the lower ionization line [S II] and the H α recombination line as compared to [O III]. While objects which have a [Ne V], He II or [Ne III] detection have a smaller scatter

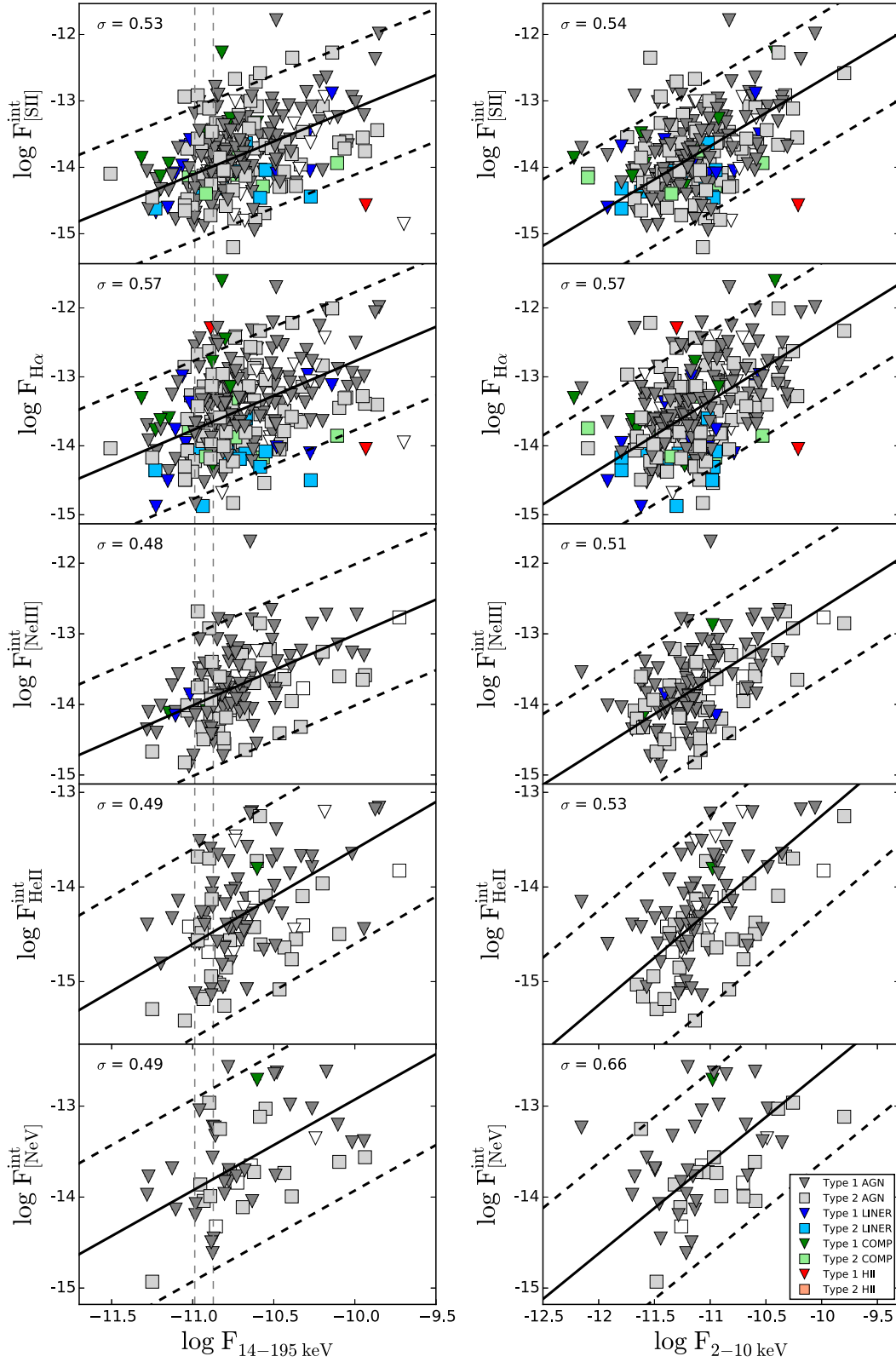


Figure 6. Plots of the relation between five optical narrow emission lines ([S II] $\lambda\lambda 6717, 6731$, H α , He II $\lambda 4686$, [Ne III] $\lambda 3869$ and [Ne V] $\lambda 3426$) with the BAT X-ray flux (left) and with the intrinsic 2–10 keV flux (right). The plots show sub-samples where the particular emission line is detected. The flux units are in $\text{erg cm}^{-2} \text{s}^{-1}$. The subplots are vertically ordered by ionization level of the emission line. The black lines are one-to-one fits of the data. The colour code represents the [N II] BPT classification. White symbols do not have any BPT classification because of undetected emission lines.

Table 2. Relations of various optical emission lines intrinsic fluxes with the 14–195 keV emission and the comparison with the [O III]–(14–195) keV relation.

Line (1)	χ [eV] (2)	N (3)	$\sigma_{[\text{O III}]}$ ^a (4)	σ_{Line} (5)	$R_{\text{Pear}, [\text{O III}]}$ ^b (6)	$R_{\text{Pear}, \text{Line}}$ (7)	p -value (8)
[O I] λ 6300	–	206	0.60 dex	0.53 dex	0.41	0.46	0.54
[S II] $\lambda\lambda$ 6717, 6731	10.4	275	0.61 dex	0.53 dex	0.37	0.38	0.89
H α	13.6	296	0.62 dex	0.57 dex	0.36	0.34	0.78
[N II] λ 6583	14.5	277	0.62 dex	0.56 dex	0.37	0.36	0.89
[Ne III] λ 3968	41.0	144	0.47 dex	0.48 dex	0.45	0.44	0.92
He II λ 4686	54.4	106	0.45 dex	0.49 dex	0.49	0.42	0.53
[Ne V] λ 3426	97.1	51	0.49 dex	0.49 dex	0.42	0.45	0.86

Note. (1) optical emission line; (2) ionization level; (3) size of the common sample in which both the Line and [O III] are detected; (4) standard deviation and (5) Pearson R coefficient of the $\log F_{[\text{O III}]}^{\text{int}} - \log F_{14-195}$ relation; (6) standard deviation and (7) Pearson R coefficient of the $\log F_{\text{Line}}^{\text{int}} - \log F_{14-195}$ relation; (8) p -value of the null hypothesis that the two correlation coefficients obtained from independent parent samples are equal.

^ascatter of the [O III]–14 195 keV relation only for objects with [O III] and the other emission line (1) detection to avoid selection biases.

^bPearson correlation coefficient for the same sub-sample used for $\sigma_{[\text{O III}]}$.

($\sigma_{[\text{O III}]}^{\text{Ne V}} = 0.49$, $\sigma_{[\text{O III}]}^{\text{Ne III}} = 0.47$, $\sigma_{[\text{O III}]}^{\text{He II}} = 0.45$), this is only because they have brighter emission lines that exclude the population of LINERs or X-ray detected AGN classified as H II/comp. In general, all lines show a statistically similar level of scatter with the X-rays independent of ionization energy.

(iii) All the optical emission lines are strongly correlated with [O III]. These correlations ($\sigma \approx 0.3$) are significantly tighter than what we find between [O III] and the X-rays.

(iv) The [O III]/X-ray flux ratio as a function of the column density (N_{H}) is consistent with being constant. However, we see a strong increase in the scatter above $N_{\text{H}} > 10^{22} \text{ cm}^{-2}$ ($\sigma_{<22} = 0.47$, $\sigma_{>22} = 0.72$). Our results show that the scatter slightly decreases towards larger physical slit sizes in kpc while the ratio $L_{[\text{O III}]}^{\text{int}}/L_{\text{X}}$ is constant for the AGN in our sample ($z > 0.01$).

We find that our study agrees well with past studies of X-ray selected AGN, but shows a weaker correlation than [O III]-selected samples which exclude X-ray selected AGN with weak emission lines (e.g. LINERs and optically elusive AGN classified as H II/comp). The correlation we find between the [O III] corrected luminosity and BAT X-ray luminosity agrees with the 9-month sample of BAT AGN (e.g. $R_{\text{Pear}}^2 = 0.41$ versus $R_{\text{Pear}}^2 \approx 0.35$ –0.40; Winter et al. 2010). Our correlation result disagrees significantly with the 9-month sample of Meléndez et al. (2008) which compared 2–10 keV X-ray fluxes from the ASCA data base which are not calculated using detailed X-ray modelling. Panessa et al. (2006) calculated the correlation between the [O III]-corrected flux and the 2–10 keV flux from a sample of 47 nearby Seyfert galaxies from the Palomar optical spectroscopic survey (Ho, Filippenko & Sargent 1995). They obtained a strong Spearman correlation coefficient of $r_s = 0.78$ that is also not in accordance with this work ($r_s = 0.39$) most likely because their sample was optically selected and did not include AGN that were X-ray weak. Our results show a slightly higher scatter than the [O III]–X-ray relation in Heckman et al. (2005) ($\sigma = 0.51$, this work: $\sigma = 0.62$) whose data come from the *RXTE* slew survey (3–20 keV).

We find no significant evidence of a weakening or strengthening of line emission at high AGN luminosities. This is because of the very weak correlation and large scatter which make a linear model inappropriate (e.g. Hogg, Bovy & Lang 2010). The improvement in the scatter of a linear fit is very small (≈ 0.01 dex) over a simple one-to-one relation between [O III] and X-ray luminosity. Using the

OLS bisector method, the best slope between the intrinsic [O III] and the 14–195 keV luminosity is 1.29 ± 0.05 and the intrinsic [O III] and 2–10 keV luminosity is 1.23 ± 0.05 . This is in accordance with previous BAT studies (e.g. 1.16 ± 0.20 ; Winter et al. 2010). Alternatively, the EM method including censored data finds a slope of 0.83 ± 0.06 which is consistent with past studies using the same line fitting method as found in Meléndez et al. 2008 (0.8 ± 0.1) or Panessa et al. 2006 (0.82 ± 0.04). In summary, our best-fitting slope is consistent with past studies, though the large intrinsic scatter of the relation yields very different results depending on the line fitting method used.

Our analysis shows that the relation between optical emission lines and hard X-rays exist, but is very weak (at best $\sigma \sim 0.5$ dex), no matter which optical emission line ([S II], H α , [O III], [Ne III], He II, [Ne V]) or which X-ray flux (2–10 keV intrinsic or 14–195 keV). Additionally, calibration, obscuration, and physical slit size play only a small role in the scatter. This scatter is much larger than the scatter of X-ray compared to SED fits of unobscured BAT AGN using the UV (e.g. 14–195 keV scatter of 0.3 dex; Vasudevan et al. 2009). While our analysis focused on the hard X-rays (> 2 keV), the soft X-rays (< 2 keV) have been shown to be dominated by emission lines in several Seyfert AGN whose extended spatial profiles show morphological correlations with the [O III] (e.g. Bianchi, Guainazzi & Chiaberge 2006; Wang et al. 2011; Koss et al. 2015). Thus, the soft X-rays coming from the emission line regions may show a tighter correlation that is worth investigating in future studies.

These results therefore strongly suggest that the majority of the intrinsic scatter has physical origins related to the NLR gas. There are several mechanisms such as covering factor of the NLR gas (e.g. Netzer & Laor 1993; Noguchi et al. 2010), density dependence of [O III] (Crenshaw & Kraemer 2005), differing ionization parameter (Baskin & Laor 2005), contamination of the narrow emission lines from star formation (e.g. Wild et al. 2011), errors in the Balmer decrement because of different distribution of Balmer lines and forbidden lines, and finally SED shape changes with luminosity (Netzer et al. 2006). Another possibility is AGN variability, where the X-ray emission varies on very short time-scales (sub-parsec scale) and represents the instantaneous accretion rate while the NLR which is orders of magnitude larger than the corona (100 parsec to kiloparsec scale), varies on much longer time-scales and traces the average accretion rate (e.g. Mushotzky, Done & Pounds 1993; Schawinski et al. 2015). Future studies using high signal-to-noise integral field

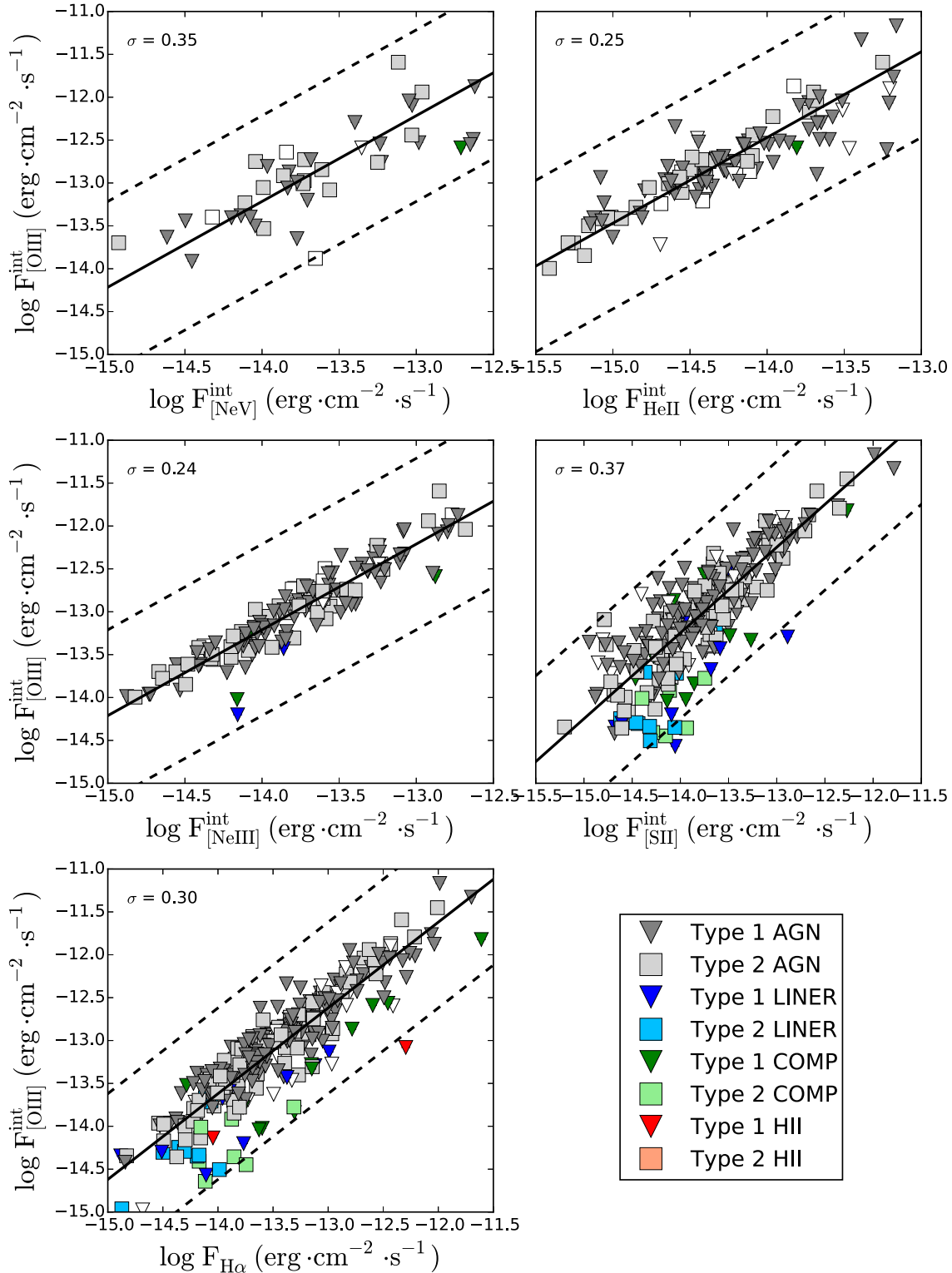


Figure 7. Plots of the relation between optical narrow emission lines and the [O III] $\lambda 5007$ line in flux space. The plots show subsamples where the particular emission line is detected. The lines are all strongly correlated with [O III]. The black lines are one-to-one fits of the data. The colour code represents the [N II] BPT classification. White symbols do not have any BPT classification because of undetected emission lines.

unit observations with wide wavelength coverage of large samples of AGN such as the Multi Unit Spectroscopic Explorer (Bacon et al. 2010) are critical to map and study the importance of these different parameters. Finally, the *Nuclear Spectroscopic Telescope Array*

(Harrison et al. 2013) hard X-ray telescope is currently observing many BAT AGN as part of the legacy survey (Baloković et al., in preparation) with a 100 times increase in sensitivity which improves estimates of X-ray luminosity compared to BAT.

Table 3. Correlation between [O III] λ 5007 and other optical emission line intrinsic fluxes.

Line (1)	χ [eV] (2)	N (3)	$\langle r \rangle$ (4)	σ (5)	R_{Pear} (6)
[O I] λ 6300	–	206	0.050	0.36 dex	0.84
[S II] $\lambda\lambda$ 6717, 6731	10.4	275	0.19	0.37 dex	0.83
H α	13.6	296	0.438	0.30 dex	0.90
[N II] λ 6583	14.5	277	0.30	0.40 dex	0.81
[Ne III] λ 3968	41.0	138	0.17	0.24 dex	0.89
He II λ 4686	54.4	106	0.031	0.25 dex	0.88
[Ne V] λ 3426	97.1	43	0.18	0.35 dex	0.79

Note. (1) optical emission line; (2) ionization level; (3) size of the sample; (4) mean flux ratio; (5) standard deviation; (6) Pearson R coefficient.

ACKNOWLEDGEMENTS

MK acknowledges support from the Swiss National Science Foundation (SNSF) through the Ambizione fellowship grant PZ00P2_154799/1. MK and KS acknowledge support from SNSF grant PP00P2_138979/1. Support for the work of ET was provided by the Center of Excellence in Astrophysics and Associated Technologies (PFB 06), by the FONDECYT regular grant 1120061 and by the CONICYT Anillo project ACT1101. The work of DS was carried out at Jet Propulsion Laboratory, California Institute of Technology, under a contract with NASA. MB acknowledges support from NASA Headquarters under the NASA Earth and Space Science Fellowship Program, grant NNX14AQ07H. CR acknowledges financial support from the CONICYT-Chile EMBIGGEN Anillo (grant ACT1101). Based on observations obtained at the Gemini Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the NSF on behalf of the Gemini partnership: the National Science Foundation (United States), the National Research Council (Canada), CONICYT (Chile), the Australian Research Council (Australia), Ministério da Ciência, Tecnologia e Inovação (Brazil) and Ministerio de Ciencia, Tecnología e Innovación Productiva (Argentina). Data in this paper were acquired through the Gemini Science Archive and processed using the Gemini IRAF package and Gemini PYTHON. Data from Gemini programmes GN-2009B-Q-114, GN-2010A-Q-35, GN-2011A-Q-81, GN-2011B-Q-96, GN-2012A-Q-28, GN-2012B-Q-25, GS-2010A-Q-54, and GS-2011B-Q80 were used in this publication. We acknowledge the work that the *Swift* BAT team has done to make this work possible. The Kitt Peak National Observatory observations were obtained using MD-TAC time as part of the thesis of MK and LW at the University of Maryland. Kitt Peak National Observatory, National Optical Astronomy Observatory, is operated by the Association of Universities for Research in Astronomy (AURA), Inc., under cooperative agreement with the National Science Foundation. Funding for SDSS-III has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, and the US Department of Energy Office of Science. This research made use of the NASA/IPAC Infrared Science Archive, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. Funding for the SDSS and SDSS-II has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, the US Department of Energy, the National Aeronautics and Space Administration, the Japanese Monbukagakusho, the Max Planck Society, and the Higher Education Funding Council for England. The SDSS website is <http://www.sdss.org/>. The

SDSS is managed by the Astrophysical Research Consortium for the Participating Institutions. The Participating Institutions are the American Museum of Natural History, Astrophysical Institute Potsdam, University of Basel, University of Cambridge, Case Western Reserve University, University of Chicago, Drexel University, Fermilab, the Institute for Advanced Study, the Japan Participation Group, Johns Hopkins University, the Joint Institute for Nuclear Astrophysics, the Kavli Institute for Particle Astrophysics and Cosmology, the Korean Scientist Group, the Chinese Academy of Sciences (LAMOST), Los Alamos National Laboratory, the Max-Planck-Institute for Astronomy (MPIA), the Max-Planck-Institute for Astrophysics (MPA), New Mexico State University, Ohio State University, University of Pittsburgh, University of Portsmouth, Princeton University, the United States Naval Observatory, and the University of Washington. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. This research made use of ASTROPY, a community-developed core PYTHON package for Astronomy (Astropy Collaboration 2013). This research made use of APLPY, an open-source plotting package for PYTHON hosted at <http://aplp.github.com>. This research has made use of the SIMBAD data base, operated at CDS, Strasbourg, France. This research made use of data and/or software provided by the High Energy Astrophysics Science Archive Research Center (HEASARC), which is a service of the Astrophysics Science Division at NASA/GSFC and the High Energy Astrophysics Division of the Smithsonian Astrophysical Observatory. This research made use of PYSPECKIT, an open-source spectral analysis and plotting package for PYTHON hosted at <http://pyspeckit.bitbucket.org>.

REFERENCES

- Alam S. et al., 2015, preprint ([arXiv:e-prints](https://arxiv.org/abs/1508.00819))
- Antonucci R., 1993, ARA&A, 31, 473
- Astropy: A community Python package for astronomy, 2013, A&A, 558, A33
- Bacon R. et al., 2010, in McLean I. S., Ramsay S. K., Takami H., eds, Proc. SPIE Conf. Ser. Vol. 7735, Ground-based and Airborne Instrumentation for Astronomy III. SPIE, Bellingham, p. 773508
- Baldwin J. A., Phillips M. M., Terlevich R., 1981, PASP, 93, 5
- Baskin A., Laor A., 2005, MNRAS, 358, 1043
- Baumgartner W. H., Tueller J., Markwardt C. B., Skinner G. K., Barthelmy S., Mushotzky R. F., Evans P. A., Gehrels N., 2013, ApJS, 207, 19
- Bianchi S., Guainazzi M., Chiaberge M., 2006, A&A, 448, 499
- Brightman M. et al., 2015, ApJ, 805, 41
- Cappellari M., Emsellem E., 2004, PASP, 116, 138
- Cardelli J. A., Clayton G. C., Mathis J. S., 1989, ApJ, 345, 245
- Crenshaw D. M., Kraemer S. B., 2005, ApJ, 625, 680
- Feigelson E. D., Nelson P. I., 1985, ApJ, 293, 192
- Ferland G. J., Osterbrock D. E., 1986, ApJ, 300, 658
- Ginsburg A., Mirocha J., 2011, Astrophysics Source Code Library, record ascl:1109.001
- Hainline K. N., Hickox R., Greene J. E., Myers A. D., Zakamska N. L., 2013, ApJ, 774, 145
- Harrison F. A. et al., 2013, ApJ, 770, 103
- Heckman T. M., Kauffmann G., Brinchmann J., Charlot S., Tremonti C., White S. D. M., 2004, ApJ, 613, 109
- Heckman T. M., Ptak A., Hornschemeier A., Kauffmann G., 2005, ApJ, 634, 161
- Ho L. C., Filippenko A. V., Sargent W. L., 1995, ApJS, 98, 477
- Hogg D. W., Bovy J., Lang D., 2010, preprint ([arxiv:1008.4686](https://arxiv.org/abs/1008.4686))
- Isobe T., Feigelson E. D., Akritas M. G., Babu G. J., 1990, ApJ, 364, 104
- Jones D. H., Read M. A., Saunders W., Colless M., Jarrett, 2009, MNRAS, 399, 683

- Kauffmann G., Heckman T. M., 2009, MNRAS, 397, 135
- Kewley L. J., Groves B., Kauffmann G., Heckman T., 2006, MNRAS, 372, 961
- Kogure T., Leung K.-C., eds., 2007, Astrophysics and Space Science Library, Vol. 342, The Astrophysics of Emission-Line Stars. Springer-Verlag, Berlin, p. 57
- Koss M. J. et al., 2015, ApJ, 807, 149
- LaMassa S. M., Heckman T. M., Ptak A., Hornschemeier A., Martins L., Sonnentrucker P., Tremonti C., 2009, ApJ, 705, 568
- Lamastra A., Bianchi S., Matt G., Perola G. C., Barcons X., Carrera F. J., 2009, A&A, 504, 73
- Lawrence A., 1991, MNRAS, 252, 586
- Massaro E., Giommi P., Leto C., Marchegiani P., Maselli A., Perri M., Piranomonte S., Sclavi S., 2009, A&A, 495, 691
- Meléndez M. et al., 2008, ApJ, 682, 94
- Mor R., Netzer H., Elitzur M., 2009, ApJ, 705, 298
- Mushotzky R. F., Done C., Pounds K. A., 1993, ARA&A, 31, 717
- Netzer H., 2013, The Physics and Evolution of Active Galactic Nuclei. Cambridge Univ. Press, Cambridge
- Netzer H., Laor A., 1993, ApJ, 404, L51
- Netzer H., Mainieri V., Rosati P., Trakhtenbrot B., 2006, A&A, 453, 525
- Noguchi K., Terashima Y., Ishino Y., Hashimoto Y., Koss M., Ueda Y., Awaki H., 2010, ApJ, 711, 144
- Ofek E. O., Oguri M., Jackson N., Inada N., Kayo I., 2007, MNRAS, 382, 412
- Panessa F., Bassani L., Cappi M., Dadina M., Barcons X., Carrera F. J., Ho L. C., Iwasawa K., 2006, A&A, 455, 173
- Revnivtsev M., Sazonov S., Jahoda K., Gilfanov M., 2004, A&A, 418, 927
- Sazonov S. Y., Revnivtsev M. G., 2004, A&A, 423, 469
- Schawinski K., Koss M., Berney S., Sartori L. F., 2015, MNRAS, 451, 2517
- Shadab A. et al., 2015, ApJS, 219, 12
- Simpson C., 1998, MNRAS, 297, L39
- Simpson C., 2005, MNRAS, 360, 565
- Steffen A. T., Strateva I., Brandt W. N., Alexander D. M., Koekemoer A. M., Lehmer B. D., Schneider D. P., Vignali C., 2006, AJ, 131, 2826
- Tueller J., Mushotzky R. F., Barthelmy S., Cannizzo J. K., Gehrels N., Markwardt C. B., Skinner G. K., Winter L. M., 2008, ApJ, 681, 113
- Ueda Y. et al., 2007, ApJ, 664, L79
- Ueda Y. et al., 2015, submitted
- Urry C. M., Padovani P., 1995, PASP, 107, 803
- Urry C. M., Maraschi L., Phinney E. S., 1991, Comments Astrophys., 15, 111
- Vasudevan R. V., Mushotzky R. F., Winter L. M., Fabian A. C., 2009, MNRAS, 399, 1553
- Veilleux S., Osterbrock D. E., 1987, ApJS, 63, 295
- Wang J. et al., 2011, ApJ, 742, 23
- Wild V. et al., 2011, MNRAS, 410, 1593
- Winter L. M., Lewis K. T., Koss M., Veilleux S., Keeney B., Mushotzky R. F., 2010, ApJ, 710, 503
- Young M., Elvis M., Risaliti G., 2010, ApJ, 708, 1388

This paper has been typeset from a \LaTeX file prepared by the author.